



Gaia-DR2 Confirms VLBA Parallaxes in Ophiuchus, Serpens, and Aquila

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Abstract

We present *Gaia*-DR2 astrometry of a sample of YSO candidates in Ophiuchus, Serpens Main, and Serpens South/W40 in the Aquila Rift, which had been mainly identified by their infrared excess with *Spitzer*. We compare the *Gaia*-DR2 parallaxes against published and new parallaxes obtained from our Very Long Baseline Array (VLBA) program Gould's Belt Distances Survey. We obtain consistent results between *Gaia* and the VLBA for the mean parallaxes in each of the regions analyzed here. We see small offsets, when comparing mean values, of a few tens of microarcseconds in the parallaxes, which are either introduced by the *Gaia* zero-point error or due to a selection effect by *Gaia* toward the brightest, less obscured stars. *Gaia*-DR2 data alone conclusively places Serpens Main and Serpens South at the same distance, as we first inferred from VLBA data alone in a previous publication. Thus, Serpens Main, Serpens South, and W40 are all part of the same complex of molecular clouds, located at a mean distance of 436 ± 9 pc. In Ophiuchus, both *Gaia* and VLBA suggest a small parallax gradient across the cloud, and the distance changes from 144.2 ± 1.3 to 138.4 ± 2.6 pc when going from L1689 to L1688.

Key words: astrometry – ISM: individual objects (Aquila Rift complex, Ophiuchus) – radiation mechanisms: non-thermal – radio continuum: stars – techniques: interferometric

1. Introduction

Observations of nearby regions rich in extremely young objects (for example Ophiuchus, Perseus, Taurus, Orion, Serpens, etc.) provide a wealth of information and clues to understanding the paradigm for star formation. In this context, the knowledge of the distance to star-forming regions is essential for properly interpreting the observed properties of young objects, from pre-stellar cores to disks, as well as their environment. An advance in constraining the distance of these star-forming regions has been made with astrometric observations using Very Long Baseline Interferometry (VLBI; see for example Reid & Honma 2014). In particular, our team has conducted a major program to measure individual distances during the last few years, via the trigonometric parallax, and proper motions to almost 100 young stars in such nearby star-forming regions (the Gould's Belt Distances Survey—GOBELINS; Loinard 2013). The distance to young stars in Ophiuchus (Ortiz-León et al. 2017b), Serpens (Ortiz-León et al. 2017a), Orion (Kounkel et al. 2017), Taurus (Galli et al. 2018), and Perseus (Ortiz-León et al. 2018) is now known with 0.2%–3% accuracy, yielding relevant information about their three-dimensional structure.

On 2018 April 25, the *Gaia* mission published its second Data Release (DR2; Gaia Collaboration et al. 2018; Lindegren et al. 2018). This release contains astrometric solutions for more than 1.3 billion stars with parallax uncertainties of ~ 0.7 mas for a magnitude $G = 20$. Several young stars within

the star-forming regions studied by our program have parallax and proper motions measurements available in the DR2 catalog. Therefore, we can now compare the VLBI astrometry against *Gaia* to validate the accuracy of our measurements.

Here we analyze *Gaia*-DR2 and Very Long Baseline Array (VLBA) astrometry in Ophiuchus, Serpens, and the Aquila Rift. Ophiuchus consists of a main cloud known as Lynds 1688 (L1688; see, e.g., Wilking et al. 2008 for a recent review) and several filamentary clouds to the northeast (L1709 and L1704) and the southeast (L1689, L1712 and L1729). The Serpens Molecular cloud and the Aquila Rift are two cloud complexes projected close to each other in the plane of the sky. The well-studied Serpens Main cluster is embedded within the Serpens Molecular cloud (Eiroa et al. 2008), while W40 (Smith et al. 1985) and the extremely young Serpens South cluster (Gutermuth et al. 2008) are embedded within the Aquila Rift, about 3° to the south of Serpens Main. Because Serpens South has the most star formation activity within the Aquila, it is often referred to as the Aquila or the Serpens-Aquila region (e.g., within the *Herschel* Gould Belt Survey program; André et al. 2010).

2. VLBA Observations

New VLBA observations toward Ophiuchus, Serpens Main, and W40 (within the Aquila Rift) were performed in the period from 2016 August to 2018 October under project code BL175. These data were taken at 5 GHz following the observing

Table 1
Astrometric Parameters from VLBA Measurements

GBS-VLA Name (1)	Other Identifier (2)	Parallax (mas) (3)	$\mu_\alpha \cos \delta$ (mas yr ⁻¹) (4)	μ_δ (mas yr ⁻¹) (5)	$a_\alpha \cos \delta$ (mas yr ⁻²) (6)	a_δ (mas yr ⁻²) (7)	Distance (pc) (8)
Ophiuchus							
J162556.09–243015.3	WLY2-11	7.181 ± 0.135	–8.098 ± 0.075	–21.978 ± 0.222	1.828 ± 0.14	4.264 ± 0.412	139.26 ± 2.62
J162557.51–243032.1	YLW24	7.089 ± 0.458	–7.047 ± 0.198	–25.135 ± 0.255	141.07 ± 9.12
J162603.01–242336.4	DoAr21	7.366 ± 0.123	–19.567 ± 0.064	–26.936 ± 0.056	135.76 ± 2.27
J162622.38–242253.3	LFAM2	7.216 ± 0.119	–5.622 ± 0.125	–26.658 ± 0.308	138.59 ± 2.28
J162629.67–241905.8	LFAM8	7.254 ± 0.112	–5.983 ± 0.059	–29.642 ± 0.129	137.86 ± 2.13
J162634.17–242328.4	S1	7.364 ± 0.052	–2.070 ± 0.010	–26.790 ± 0.021	135.79 ± 0.97
J162642.44–242626.1	LFAM15	7.176 ± 0.058	–6.327 ± 0.016	–26.905 ± 0.036	139.36 ± 1.13
J162643.76–241633.4	VSSG11	7.160 ± 0.152	–10.48 ± 0.16	–38.99 ± 0.35	0.31 ± 0.65	–1.54 ± 1.09	139.7 ± 3.0
J162649.23–242003.3	LFAM18	7.176 ± 0.055	–8.627 ± 0.159	–20.015 ± 1.064	139.35 ± 1.06
J162718.17–242852.9	YLW 12Bab	7.135 ± 0.064	10.875 ± 0.049	–25.057 ± 0.140	–0.622 ± 0.116	–0.177 ± 0.018	140.16 ± 1.26
J162718.17–242852.9	YLW 12Bc ^a	7.135 ± 0.064	–11.219 ± 0.052	–23.107 ± 0.128	0.195 ± 0.027	0.109 ± 0.066	140.16 ± 1.26
J162721.81–244335.9	ROXN39	7.396 ± 0.071	–6.858 ± 0.095	–24.864 ± 0.146	135.22 ± 1.30
J162721.97–242940.0	GY256	6.711 ± 0.111	–6.673 ± 0.054	–34.41 ± 0.176	149.02 ± 2.47
J162726.90–244050.8	YLW15	7.376 ± 0.095	–12.841 ± 0.042	–26.255 ± 0.169	135.57 ± 1.75
J162730.82–244727.2	DROXO71	7.455 ± 0.229	–4.822 ± 0.144	–28.256 ± 0.253	134.14 ± 4.12
J162804.65–243456.6	ROXN 78	7.185 ± 0.091	–5.446 ± 0.028	–29.301 ± 0.071	139.18 ± 1.77
J163035.63–243418.9	SFAM 87	7.216 ± 0.068	–7.702 ± 0.019	–26.028 ± 0.031	138.58 ± 1.31
J163115.01–243243.9	ROX42B	6.922 ± 0.043	–5.817 ± 0.033	–23.206 ± 0.259	–0.506 ± 0.041	0.962 ± 0.315	144.47 ± 0.91
J163151.93–245617.4	...	7.265 ± 0.778	–8.564 ± 0.268	–27.132 ± 0.589	138 ⁺¹⁷ _{–13}
J163152.10–245615.7	LDN1689IRS5	6.677 ± 0.157	–6.537 ± 0.09	–22.557 ± 0.144	149.76 ± 3.52
J163200.97–245643.3	WLY2-67	6.741 ± 0.173	–5.699 ± 0.14	–23.994 ± 0.357	148.34 ± 3.80
J163211.79–244021.8	DoAr51	6.972 ± 0.041	–4.746 ± 0.084	–23.139 ± 0.099	143.43 ± 0.85
Serpens Main							
J182933.07+011716.3	GFM 11	2.313 ± 0.078	3.634 ± 0.050	–8.864 ± 0.127	432 ⁺¹⁵ _{–14}
J182957.89+011246.0	EC 95	2.307 ± 0.022	3.579 ± 0.021	–8.359 ± 0.023	433 ± 4
J183000.65+011340.0	GFM 65	2.375 ± 0.222	2.437 ± 0.357	–8.263 ± 0.228	421 ⁺⁴⁴ _{–36}
W40							
J183114.82–020350.1	KGF 36	2.297 ± 0.116	0.400 ± 0.065	–6.607 ± 0.063	435 ⁺²³ _{–21}
J183123.62–020535.8	KGF 97	2.119 ± 0.077	–0.300 ± 0.047	–7.432 ± 0.048	472 ⁺¹⁸ _{–17}
J183126.02–020517.0	KGF 122	2.261 ± 0.138	1.845 ± 0.645	–6.029 ± 0.339	442 ⁺²⁹ _{–25}
J183127.45–020512.0	KGF 133	2.194 ± 0.231	0.181 ± 0.148	–8.750 ± 0.323	0.270 ± 0.158	–0.53 ± 0.344	456 ⁺⁵⁴ _{–44}
J183127.65–020509.7	KGF 138	2.353 ± 0.106	0.172 ± 0.089	–6.784 ± 0.785	425 ⁺²⁰ _{–18}

Note.

^a Parallax is fixed at the value obtained for YLW 12Bab when solving for the other astrometric parameters.

strategy described in detail in Ortiz-León et al. (2017a, 2017b, hereafter Papers I and II, respectively).

The resulting astrometric parameters from the updated fits, including new sources, are given in Table 1. In addition to the 24 objects published in Papers I and II, here we present new parallaxes for six objects. Following the approach outlined in Loinard et al. (2007), the data are fitted with a model that assumes a uniform proper motion and has the following free parameters: parallax (ϖ), proper motions (μ_α , μ_δ), and position at the reference epoch (α_0 , δ_0). The fits to sources that are members of wide binary systems also include acceleration terms, (a_α , a_δ). We also update the orbital fits of our binary systems following Kounkel et al. (2017); however, in this Letter we will focus solely on parallax and leave the discussion on the orbital parameters for a forthcoming publication.

For Ophiuchus, the weighted mean of individual parallaxes in each cloud gives $\varpi = 7.23 \pm 0.14$ mas and $\varpi = 6.93 \pm 0.06$ mas, for L1688 and L1689, respectively. Herein, and in the rest of this Letter, the quoted errors on weighted mean

parallaxes correspond to the standard deviation, unless otherwise noted. These measurements suggest that, although small, there is a parallax gradient across this cloud, but the difference in parallax is only significant at 2σ .

In Serpens/Aquila, we found $\varpi = 2.31 \pm 0.01$ mas and $\varpi = 2.23 \pm 0.10$ mas, for the Serpens Main and W40 clusters, respectively. In this case, the mean parallaxes are consistent between them within 1σ . The average of all sources in both regions gives $\varpi = 2.30 \pm 0.05$ mas.

3. Analysis of *Gaia* DR2 Data

We use the list of young stellar object (YSO) candidates derived with *Spitzer* by Dunham et al. (2015) for Ophiuchus (292 sources), Serpens (227 sources), and Aquila (1319 sources). However, this sample suffers from contamination by background AGB stars (between 25% and 90%; Dunham et al. 2015), so their membership and nature as true YSOs must be confirmed. We also use the list of 316 YSOs compiled by Wilking et al. (2008) for Ophiuchus, which are all 2MASS

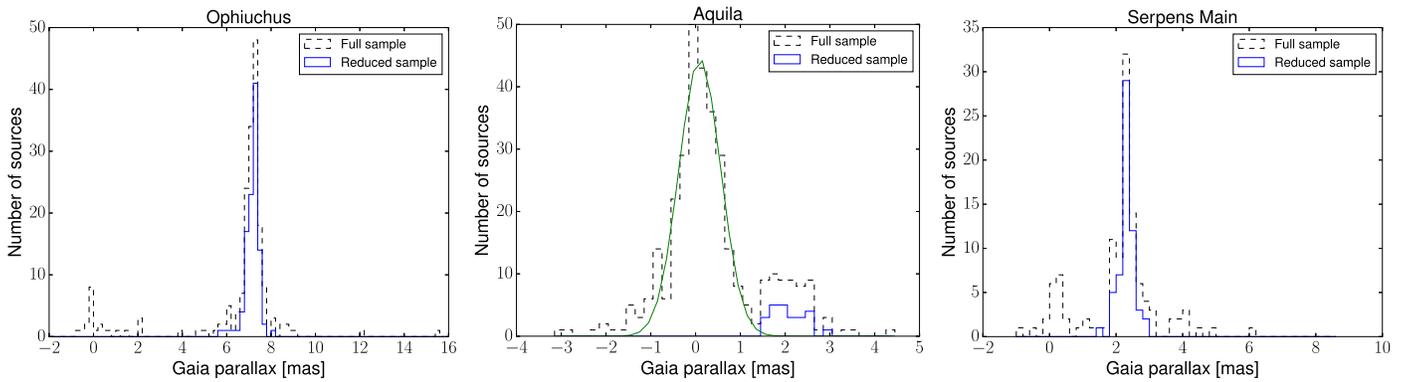


Figure 1. *Gaia* parallaxes from DR2 in Ophiuchus (left), Aquila (middle), and Serpens Main (right). The blue histograms show the distribution of parallaxes after cutting stars according to the criteria given in the text. The green line in the middle panel is a Gaussian fit to the parallax distribution with $\varpi \lesssim 1.4$ mas.

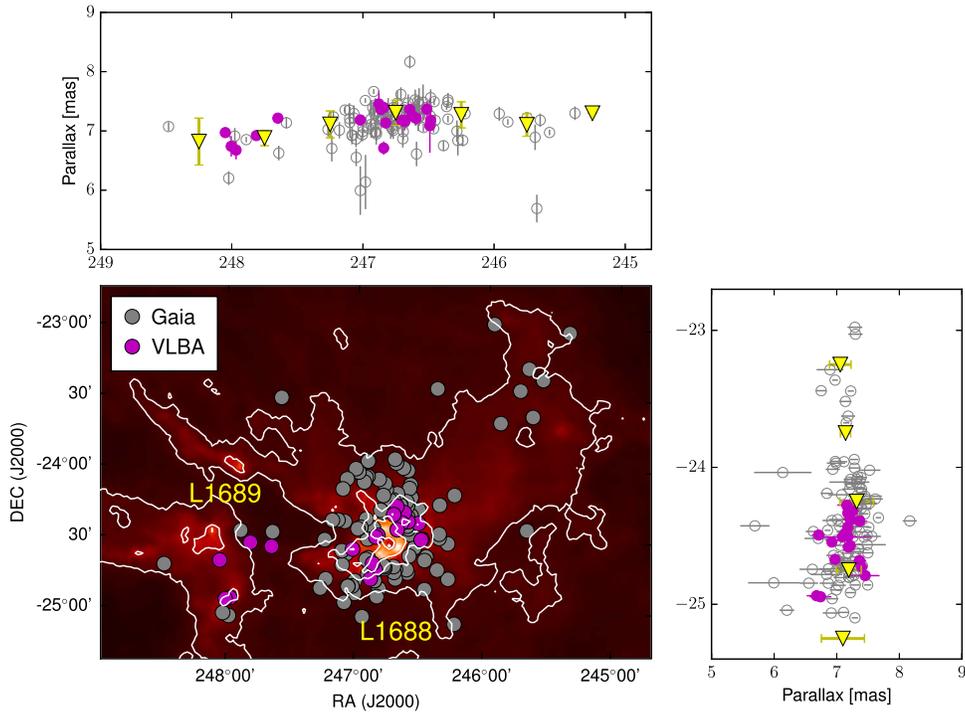


Figure 2. *Gaia*-DR2 and VLBA parallaxes measured in Ophiuchus. The inner panel shows the spatial distribution of stars with measured parallaxes, while the top and right panels show the parallaxes as a function of R.A. and decl., respectively. The gray open circles are stars with *Gaia* parallaxes, while the magenta filled circles are stars with VLBA parallaxes. The yellow triangles are the weighted mean of *Gaia* parallaxes determined for bins with a width of $0.5''$. The map in the background is an extinction map obtained as part of the COMPLETE project (Ridge et al. 2006). The white contours indicate $A_V = 4, 12,$ and 20 .

sources within the L1688 cloud, with detection in the K_S band and other signs of youth, such as H_α or X-ray emission. The combined sample has 565 YSO candidates in Ophiuchus.

We cross-matched the positions of the YSO candidates against *Gaia* DR2 positions using a search radius of $1''$ and the Virtual Observatory tool TOPCAT (Taylor 2005). In Ophiuchus, 191 sources of the *Spitzer*+2MASS catalog have astrometric solutions. The distribution of *Gaia* parallaxes is shown in Figure 1 (left panel), and their positions on the plane of the sky are shown in Figure 2. We note that this sample contains background and foreground contaminants with parallaxes $\lesssim 3$ and $\gtrsim 11$ mas (i.e., well outside the range from 120 to 150 pc reported in the literature as the distance to Ophiuchus). To remove these stars, we look at the distributions of proper motions and, following Ortiz-León et al. (2018) and Dzib et al. (2018), we cut all stars with proper motions that are beyond 3σ from the mean, as determined

by fitting a Gaussian model to the measured proper motions in each direction. Then, we produce a sample with reliable parallaxes by applying the recommended cuts given in Lindegren et al. (2018), which are designed to remove sources with poor or spurious astrometric solutions. Specifically, we applied the criteria expressed in Equations (C.1) and (C.2) in Appendix C of Lindegren et al. (2018). After also cutting sources that lie well outside the region shown in Figure 2 (two sources at $(\alpha, \delta) \sim (252^\circ, -14^\circ)$), the sample is reduced to 107 objects. The resulting weighted mean parallax for this reduced sample is 7.23 ± 0.25 mas.

To compare against the VLBA parallaxes, we bin the *Gaia* parallaxes into R.A. bins of width equal to $0.5''$ and take the weighted average of the parallaxes within each bin. The resulting values are plotted in Figure 2 as yellow triangles. The weighted parallaxes for the two first bins are 6.82 ± 0.40 and 6.88 ± 0.13 mas, respectively, which agree within the

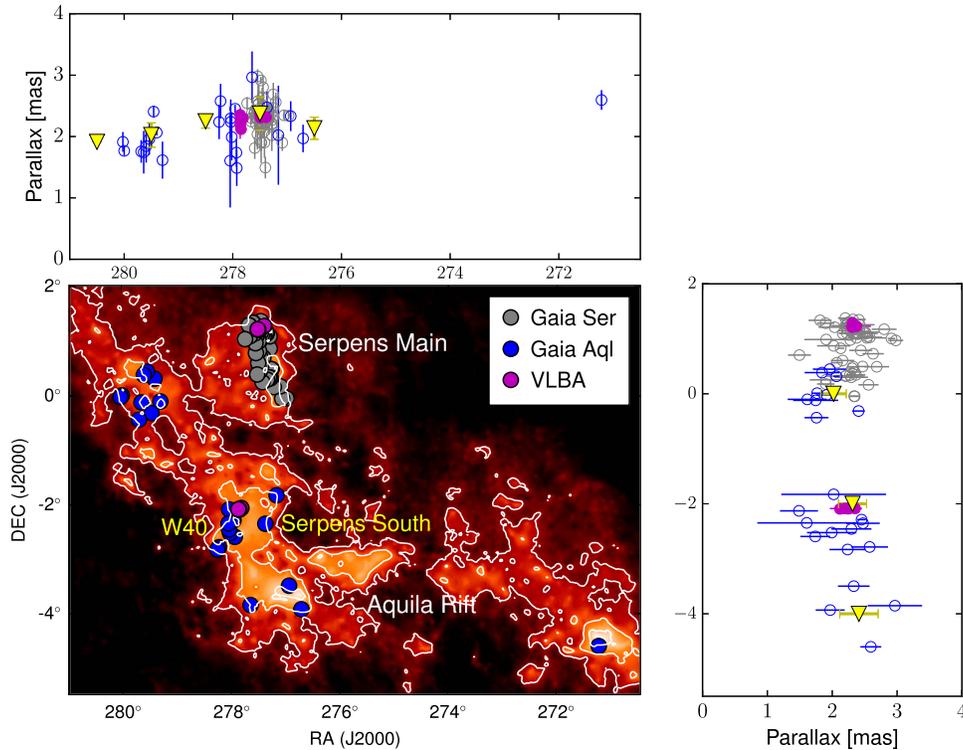


Figure 3. Same as Figure 2 but for the Serpens/Aquila region. Here the bins have widths of 1° and 2° , in R.A. and decl., respectively. The yellow triangles are the weighted mean of *Gaia* parallaxes in Aquila determined for these bins. *Gaia* parallaxes measured in Aquila are shown as blue circles. The extinction map in the background was taken from Cambrésy (1999). The white contours indicate $A_V = 4, 7,$ and 9 .

errors with the VLBA parallax derived for L1689 (6.93 ± 0.06 mas). *Gaia* data also suggest that the parallaxes at the center of L1688 are, on average, slightly larger than L1689. The weighted average parallax rises to 7.31 ± 0.21 mas roughly at the center of L1688.

For Aquila, we found 341 YSO candidates with available parallaxes in DR2. We can see in the middle panel of Figure 1 that the number of background contaminants is remarkably high for this region ($\sim 80\%$ of the total), and the presence of two distributions, with peaks separated at $\varpi \sim 2$ mas, is clear. The stars with small parallaxes almost follow a Gaussian distribution as can be seen in the same figure. The center of this Gaussian is at $\varpi = 0.09 \pm 0.03$ mas and the width is 0.46 mas. From this figure, we can assume that all stars with $\varpi \lesssim 1.4$ mas, which is the boundary between the Gaussian and the second distribution, are background contaminants. We then further cut the sample by applying the proper motion criteria and the criteria to filter stars with spurious astrometric solutions as in Ophiuchus. This produced the reduced sample of 24 stars shown in Figure 3, where we see that stars are grouped in at least three regions. One group of stars is distributed within a diameter of $\sim 2^\circ$ around the position $(\alpha, \delta) \sim (277^\circ.6, -2^\circ.3)$, encompassing the W40 and Serpens South clusters. A second group of stars is located around the position $(\alpha, \delta) \sim (279^\circ, +0^\circ)$. The next group is centered at $(\alpha, \delta) \sim (277^\circ, -4^\circ)$, and finally one more star is close to $(\alpha, \delta) \sim (271^\circ, -4^\circ.5)$.

Similarly to Ophiuchus, we bin the *Gaia* parallaxes into R.A. and decl. bins of width equal to 1° and 2° , respectively. We then take the weighted average in each bin (the yellow triangles in Figure 3). The groups identified above in the spatial distribution of the *Gaia* parallaxes are clearly seen in the plot corresponding to the decl. direction (right panel in Figure 3). The weighted mean of stars within a radius of 1° around

$(\alpha, \delta) \sim (277^\circ.6, -2^\circ.3)$, i.e., the area covering the W40 and Serpens South clusters and corresponding to the second yellow triangle in the right panel of Figure 3, gives $\varpi = 2.31 \pm 0.22$ mas. This is the representative *Gaia* value that we take for W40/Serpens South.

In Serpens, only 68 *Spitzer* sources have parallaxes available in DR2. We also cross-matched the catalog of optical candidate young stars published by Erickson et al. (2015), which contains 62 candidate members based on the presence of H α emission, lithium absorption, X-ray emission, mid-infrared excess, and/or reflection nebulosity. For this catalog, we use a match radius of $2''$ to be more conservative. This adds 47 stars to the number of sources with *Gaia* parallaxes in Serpens. The parallax distribution is shown in the right panel of Figure 1. Again, we see a peak in the parallax distribution due to background contaminants, which we remove by using the proper motion criteria. The weighted mean parallax for Serpens (of the 59 stars left after cleaning the sample from spurious astrometric solutions) gives $\varpi = 2.32 \pm 0.18$ mas. We take this value as the representative *Gaia* parallax for Serpens Main.

4. Discussion

The *Gaia*-DR2 measurements confirm the results reported in Papers I and II, which were based on VLBA data alone. In Paper II, we derived parallaxes for stars in W40 and Serpens Main. The VLBA parallaxes for both regions were consistent between them and suggested a common distance of ~ 436 pc. We argued that, given that Serpens South is projected very close to the W40 cluster, it should be located at the same distance of ~ 436 pc as W40, which implied that the three regions, i.e., Serpens Main, Serpens South, and W40, are part of the same large complex of molecular clouds. This

conclusion was supported by the similar local standard of rest velocities obtained from observations of isotopologues of CO toward IRAS 18275–0203 and a nearby embedded source in Serpens South (Gutermuth et al. 2008) and toward various positions within the Serpens Molecular cloud (White et al. 1995).

The values of the mean parallax derived from the *Gaia*-DR2 data alone for Serpens South/W40 and Serpens Main ($\varpi = 2.31 \pm 0.22$ mas and $\varpi = 2.32 \pm 0.18$ mas, respectively) are identical within the errors, which conclusively places both regions at a similar distance (432_{-37}^{+45} pc and 430_{-30}^{+36} pc, respectively), i.e., at a significantly larger distance than the old, but still used by some authors, value of 260 pc (Straizys et al. 2003). As seen in the third panel of Figure 1, there is a group of stars that have smaller distances ($\varpi \gtrsim 4$ mas or $d \lesssim 250$ pc), but they represent a minority compared to the main body of young stars, which draw a clear distribution around 2.3 mas.

When compared against the VLBA parallaxes, we see an offset of $\varpi_{Gaia} - \varpi_{VLBA} = +0.09 \pm 0.07$ mas and $\varpi_{Gaia} - \varpi_{VLBA} = +0.02 \pm 0.02$ mas, for W40/Serpens South and Serpens Main, respectively, where ϖ_{Gaia} and ϖ_{VLBA} are the weighted mean parallax for *Gaia* and the VLBA, respectively. In Ophiuchus, we measure offsets of $\varpi_{Gaia} - \varpi_{VLBA} = +0.06 \pm 0.02$ mas and $\varpi_{Gaia} - \varpi_{VLBA} = -0.07 \pm 0.04$ mas, in L1688 and L1689, respectively. Although these offsets are relatively small, it is important to understand their origin. The offset measured for stars in L1689 is consistent with the zero-point error of *Gaia* DR2 parallaxes, which has a mean of -0.029 mas (Lindgren et al. 2018). The *Gaia*-VLBA parallax offset is positive in L1688, Serpens Main, and the Aquila, which means that the *Gaia* parallax there is larger than the VLBA parallax, or that *Gaia* distances are, on average, smaller than the VLBA distances. This is actually the behavior we expect to see in regions where the optical extinction is high ($A_V \gtrsim 20$ as is seen in Serpens and Ophiuchus; Ridge et al. 2006; Bontemps et al. 2010). Since *Gaia* observes in the optical, it can only detect stars that are not too obscured by dust extinction, while the VLBA can see through the extinction wall. Thus, *Gaia* could be biased toward detecting stars preferentially on the near side of the clouds, whereas the VLBA does not suffer from such a bias. This effect is not evident in L1689 or IC 348 in Perseus (Ortiz-León et al. 2018), where the optical extinction is lower ($A_V \lesssim 10$; Ridge et al. 2006—see also Kounkel et al. 2018 for a parallax comparison of the entire YSOs sample in common with *Gaia* and GOBELINS).

Thus, our findings suggest *Gaia* would be biased in regions with strong optical extinction. We were not able to investigate the magnitude of the bias as a function of extinction from ancillary data. Visual extinction maps available in the literature have been derived from different data sets, using different algorithms. For instance, the Ophiuchus map shown in Figure 2 was obtained from 2MASS near-infrared data, while the Serpens map (Figure 3) was derived from optical star counts and is only tracing the first layer of the extinction wall. From the distribution of the extinction measured in Ophiuchus, which increases from the east to the core of the cloud, we suggest that *Gaia* parallaxes on regions with $A_V \gtrsim 18$ may show a significant bias.

It is also important to note that the error bars on *Gaia* parallaxes are on average significantly larger than the errors on VLBA parallaxes, mainly in the Aquila Region, and that the

individual *Gaia* values show a larger dispersion than the dispersion seen for VLBA sources. This is clearly seen in Figures 2 and 3. Quantitatively, this comes from taking the mean of the *Gaia* parallax errors of the samples in each region before cleaning from possible poor astrometric solutions. The mean errors are $\sigma_{\varpi, Gaia} = 0.24, 0.18,$ and 0.40 mas in Ophiuchus, Serpens, and Aquila, respectively, while the mean of VLBA parallax errors is $\sigma_{\varpi, VLBA} = 0.12$ mas in the three regions.

Given that the *Gaia* zero-point offset applicable to a given region in the sky is still not well determined (Lindgren et al. 2018), and the corrections to be applied to the *Gaia* parallaxes are still poorly constrained, we do not attempt here to correct the *Gaia* parallaxes. We instead recommend the use of the VLBA parallaxes and distances, as they still represent the most accurate measurements obtained so far toward these three regions. Inverting the VLBA parallaxes yields distances of 138.4 ± 2.6 pc, 144.2 ± 1.3 pc, and 436 ± 9 pc for L1688, L1689, and Serpens/Aquila, respectively.

Regarding the presence of parallax gradients across the clouds, we do see in Figures 2 and 3 that *Gaia* parallaxes change with position, rising toward the central part of the clouds, which is expected if extinction is biasing more toward the center. In Ophiuchus, VLBA parallaxes show clear variations with position along the R.A. direction, which supports the idea that different structures in the cloud are located at slightly different distances. In Serpens/Aquila, the significance of such parallax variations is low. However, we do not rule out a possible difference of ~ 16 pc between Serpens Main and Serpens South/Aquila.

To strengthen our conclusion on the distance to Aquila, we use all stars from the *Gaia* DR2 catalog covering a circular area with a radius of 3° and centered at $(\alpha, \delta) = (277^\circ.5, -1^\circ.0)$ (i.e., at the mid-point between Serpens Main and Serpens South). We only keep stars with $\varpi > 1$ mas ($d < 1$ kpc) and $\sigma_{\varpi} < 0.5$ mas, resulting in a total of 79550 stars. The spatial distribution of these stars is shown in Figure 4, where each point represents the average of all stars within small patches of size, set to 500 arcsec, whose average distance (estimated as $1/\overline{\varpi}$) is indicated by a color code. We see in Figure 4 the effect of the obscuration by the cloud, because, on average, the areas with large extinction are the ones that have the lower distances. The distribution of distances is also shown in Figure 4. Here we see that the number of sources continues to rise with distance until it reaches ~ 430 pc, where there is a pause in the rise. This distance corresponds to the main distance to the cloud, which agrees very well with the distance derived from the VLBA parallax measurements.

The nature of the background *Spitzer* contaminants in Aquila is still unknown, but they are likely AGB stars according to Dunham et al. (2015). The proper motions of these stars are nonzero, which suggests they are Galactic sources.

5. Conclusions

We used *Gaia*-DR2 and VLBA data to investigate the distance to Ophiuchus, Serpens Main, and W40/Serpens South in the Aquila Rift region. Our target samples for the *Gaia* analysis consisted of YSO candidates, with infrared excess identified by *Spitzer*, which were complemented by including 2MASS and optical sources having other signs of an association to young stars. We carefully cleaned our samples from background and foreground stars, and from stars with spurious astrometric

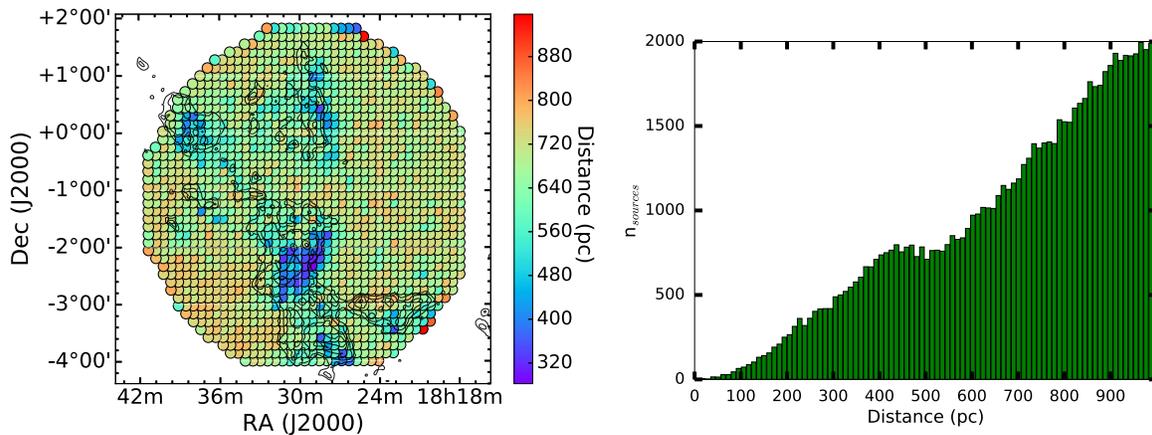


Figure 4. Left: *Gaia*-DR2 distances for all stars within 3° around $(\alpha, \delta) = (277.5, -1.0)$ and with $\varpi > 1$ mas and $\sigma_\varpi < 0.5$ mas, overlaid on the extinction map of the Aquila. Each point is the average of all stars within small patches of size set to 500 arcsec. Right: histogram of distances for the same sample of stars.

solutions, and then derived the weighted mean parallax of each region. From *Gaia* data alone, we found identical parallaxes for both Serpens Main ($\varpi_{\text{Gaia}} = 2.32 \pm 0.18$ mas) and W40/Serpens South ($\varpi_{\text{Gaia}} = 2.31 \pm 0.22$ mas), which are consistent with the parallaxes measured independently by the VLBA ($\varpi_{\text{VLBA}} = 2.30 \pm 0.05$ mas). Hence, this confirms that the three regions are part of the same complex of molecular clouds. VLBA and *Gaia*-DR2 parallaxes are also highly consistent between them in Ophiuchus, where we found $\varpi_{\text{L1688, VLBA}} = 7.23 \pm 0.14$ mas, $\varpi_{\text{L1689, VLBA}} = 6.93 \pm 0.06$ mas, $\varpi_{\text{L1688, Gaia}} = 7.29 \pm 0.22$ mas, and $\varpi_{\text{L1689, Gaia}} = 6.86 \pm 0.23$ mas. The small offsets between *Gaia* and the VLBA can be understood in terms of the *Gaia* parallax zero-point error or *Gaia* being biased toward the brightest and less obscured stars.

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